



## Measurement of the cosmic electron plus positron spectrum with the MAGIC telescopes

D. BORLA TRIDON<sup>1</sup>, P. COLIN<sup>1</sup>, L. COSSIO<sup>2</sup>, M. DORO<sup>3</sup>, V. SCALZOTTO<sup>4</sup>,

ON BEHALF OF THE MAGIC COLLABORATION

<sup>1</sup>*Max-Planck-Institut für Physik, Munich, Germany*

<sup>2</sup>*Università di Udine and INFN sez. di Trieste*

<sup>3</sup>*Universitat Autònoma de Barcelona, Barcelona, Spain*

<sup>4</sup>*Università di Padova and INFN sez. di Padova*

*dborla@mppmu.mpg.de*

**Abstract:** Cosmic electrons with energies in the TeV range lose their energy rapidly through synchrotron radiation and inverse Compton processes, resulting in a relatively short lifetime ( $\sim 10^5$  years). They are only visible from comparatively nearby sources ( $< 1$  kpc). Unexpected features in their spectrum at a few hundreds GeV, as measured by several experiments (ATIC, Fermi and H.E.S.S. among others), might be caused by local sources such as pulsars or by dark matter annihilation/decay. In order to investigate these possibilities, new measurements in the TeV energy region are needed. Since the completion of the stereo system, the MAGIC Cherenkov experiment is sensitive enough to measure the cosmic electron flux between a few hundred GeV and few TeV. The electron signal has to be extracted from the overwhelming background of hadronic cosmic rays estimated through Monte Carlo simulations. Here we present the first results of the cosmic electron spectrum measured with the MAGIC telescopes.

**Keywords:** TeV cosmic electrons, pulsar, dark matter, Cherenkov telescopes

## 1 Introduction

TeV cosmic electrons and positrons (hereafter electrons) lose energy mainly due to inverse Compton processes and due to synchrotron radiation caused by the weak cosmic magnetic field. These losses limit the distance at which they can be observed. Therefore, their energy spectrum can give us clues about their origin, providing useful information on the nearby Universe.

Recently, several measurements of cosmic electrons have shown some features in their spectra, which have excited astrophysicists in the field. Many interpretations in terms of dark matter scenarios [1] or astrophysical sources such as pulsars [2] or supernova remnants [3] are invoked.

Until recently, energy spectra measurements of electrons were obtained by balloon and satellite experiments. At TeV energies, however, the sensitivity of these instruments is insufficient due to their small sizes and short exposure time of flight. Nowadays, ground-based Cherenkov telescopes, with their large collection areas and good sensitivity, represent an excellent tool for measuring high energy cosmic electrons, via the indirect observation of the air showers that charged particles generate in the atmosphere. In 2009 H.E.S.S. measured the electron spectrum from 300 GeV up to  $\sim 4$  TeV [7]. MAGIC, with its two largest-dish Cherenkov telescopes world-wide, is now one of the most suited experiments that can contribute

to this measurement. It has the potential to overlap with the energy range of the other experiments (Fermi [4], ATIC [5] and H.E.S.S. [6] [7] among the others), confirming previous measurements of the electron spectrum to TeV energies.

The determination of the cosmic electron spectrum is pursued through a non-standard analysis, which is instead optimized for the reconstruction of images from  $\gamma$ -ray showers. Nonetheless, as in the case of  $\gamma$ -rays, electrons events are largely subdominant, and overwhelmed by a much larger background of hadronic events and also  $\gamma$ -ray events. While  $\gamma$ -rays are not deflected by the magnetic fields and thus their arrival direction points directly to their source, electrons are isotropically diffused. Therefore, the identification of charged particles cannot be done through the arrival direction information, but only via the shape of the image, which is produced by the air showers in the atmosphere and recorded by the telescopes. The data are modelled with simulations of electron showers and the background is rejected by applying selection criteria. Since this method does not separate electrons from gammas, a small contamination from diffuse gammas is expected. Moreover, Cherenkov telescopes do not separate the charge of the particles, thus the presented measurement includes the contribution of both electrons and positron.

## 2 The MAGIC telescopes

MAGIC (*Major Atmospheric Gamma Imaging Cherenkov*) is a stereoscopic system of two imaging atmospheric Cherenkov telescopes (IACT). Among the operational IACTs, it has the world-largest dishes with 17m diameter each. They are located at the Observatory of the Roque de los Muchachos on the Canary Island of La Palma (28.75° N, 17.86° W, 2200 m a.s.l.).

Since fall 2009 MAGIC is fully operational in stereoscopic mode. The two telescopes are designed to detect very high energy (VHE)  $\gamma$ -rays in the energy range from 50 GeV to tens of TeV and can also detect electrons in the same energy range.

## 3 Analysis

The data used in this analysis come from observations of selected extragalactic sky areas (in which no  $\gamma$ -ray emission from astrophysical objects has been found) in order to minimize the contribution from the diffuse  $\gamma$ -ray emission which comes mainly from the Galactic plane. A priori, a search for a gamma-point source has been carried out to demonstrate that no gammas contaminated the sample. The data were extracted from observation carried out in December 2009, June 2010, October 2010 and November 2010. Data passing quality selection criteria, with zenith angles between 14° and 27°, have been used in the analysis, providing a total amount of  $\sim 14$  hours. The images are cleaned using a threshold of 6 (9) photoelectrons (phes) for core pixels and 3 (4.5) phes for boundary pixels has been applied for the MAGIC-I (MAGIC-II) telescope data respectively. Higher cleaning for the MAGIC-II telescope data is needed because of a higher photon to photoelectron efficiency and a higher noise level in the read-out chain compared to MAGIC-I read-out. Beside, the arrival time in each pixel belonging to shower image can deviate at most of 4.5 ns from the shower core arrival time. Therefore we set a maximum time difference between adjacent pixels to be less than 1.5 ns.

The most important issue in the electron analysis is the electron/hadron separation and the rejection of the hadronic background. A Random Forest (RF) method [9] is used to compute the *Hadronness* parameter, which spans from 0 to 1 and gives the degree of hadron-like. A classification tree of the RF is trained with a sample of Monte Carlo (MC) electrons and a sample of MC protons. The input parameters of the RF are the Hillas parameters [8] of the shower images (*Size*, *Width*, *Length* for both the telescopes), the *Impact* parameter of the two telescopes and the reconstructed *Height* of the shower maximum. A cut on the number of *Islands* parameter has been applied for both the telescopes.

The reconstruction of the energy of each event is

done via a look-up-table obtained also from MC events, based on *Size*, *Impact*, *Height* and *Zenith* angle. The mean energy resolution is below 20 % in the energy range between 100 GeV and 2 TeV.

The signal of the diffuse cosmic electrons is determined in the phase space of the *Hadronness* parameter (the result of the RF). In the distribution of the *Hadronness* the diffuse electron signal peaks around *Hadronness*=0, while the hadronic background around *Hadronness*=1. In order to determine the electron spectrum, the following procedure is applied:

1. Apply selection cuts.
2. Determine the *Hadronness* distribution for data and background, define a signal region and normalize the two distributions by the number of events in an optimized non signal region.
3. Count events in the signal region ( $N_{on}$ ).
4. Count background events in the signal region ( $N_{bg}$ ).
5. Determine the number of excess in the signal region ( $N_{excess} = N_{on} - N_{bg}$ ).
6. Determine the effective observation time, effective acceptance and finally the energy spectrum.

Selection cuts are applied in order to rejected part of the background. The *Impact* is selected to be within 10 and 300 m referred to both telescopes, for MAGIC-I *Size* > 100 phes and MAGIC-II *Size* > 200 phes. The signal region, by means of a cut in the *Hadronness* parameter, is chosen by requiring an acceptance for MC electrons of 60%. The background distribution is determined through MC proton simulations. We underline that by now we limited ourselves in the production of only proton events. This choice is motivated by the fact that protons constitute by far the major component of the total CR spectrum (followed by Helium). On the other hand, we mention that a faulty description of the background may alter our results consistently. This is made even more complicate by the fact that hadronic interactions are difficult to simulate with high precision: many types of interactions compete, and not all details of cross sections are known, particularly at high energies [10]. The hadronic showers have been simulated with CORSIKA [11], using the FLUKA model [12] for the low energies, in combination with the QGSJet-II [13] interaction model for the high energies. MC proton events have been simulated in a *Zenith* angle range between 5° and 30° with a maximum *Impact* parameter of 1.2 km in a solid angle of 0.034 sr. Since the MC proton simulations have been performed with a different energy spectrum<sup>1</sup> of  $E^{\Gamma_{sim}}$  (with  $\Gamma_{sim} = -2.0$  or  $-1.78$ ) compared to the real cosmic-ray spectrum of  $E^{\Gamma_{real}}$ , the *Hadronness* distribution has to be corrected. Thus, to each event  $i$  of the *Hadronness* distribution a weight factor of  $w_i = E_{true}^{(\Gamma_{real} - \Gamma_{sim})}$  is assigned;

1. to increase the statistics at high energy.

$E_{true}$  is the simulated energy. In this case, while the uncertainties on the ON data are computed according to the Poisson statistic, the uncertainties on the MC protons are defined as  $\Delta N_{pr} = \sqrt{\sum_i w_i^2}$ , the square root of the sum of the weights in the considered bins of the *Hadronness* distribution. The normalization factor  $\alpha = \frac{n_{on}}{n_{pr}}$  is the ratio between the numbers of ON and MC proton events in the non-signal region which implies  $N_{bg} = \alpha N_{pr}$ . The significance of the excess is defined as:

$$S = \frac{N_{excess}}{\Delta N_{bg}} = \frac{N_{on} - \alpha N_{pr}}{\alpha \left[ \Delta N_{pr}^2 + N_{pr}^2 \left( \frac{1}{n_{on}} + \frac{\Delta n_{pr}^2}{n_{pr}^2} \right) \right]^{1/2}}$$

The effective acceptance of the telescope system, used for the spectrum calculations, is calculated after the solid angle of acceptance of the simulated electrons, which in our case corresponds to  $\Omega = 0.019$  sr. The electrons are simulated in a circular area of 650 m radius in a *Zenith* angle range between  $5^\circ$  and  $30^\circ$ .

## 4 Results

The total *Hadronness* distribution of the observed events is plotted together with the normalized distribution of background events in figure 1. The normalization is done in the *Hadronness* range between 0.4 and 0.8. In total, 5383 electron excess events are seen with a significance of  $8.08\sigma$  in the energy range between 150 GeV and 2 TeV. In figure 2 the *Hadronness* distribution of the excess events is instead compared with that of the MC electron simulations, at low *Hadronness* values, where the electron signal is expected. The two distributions are in agreement, demonstrating the electron excess.

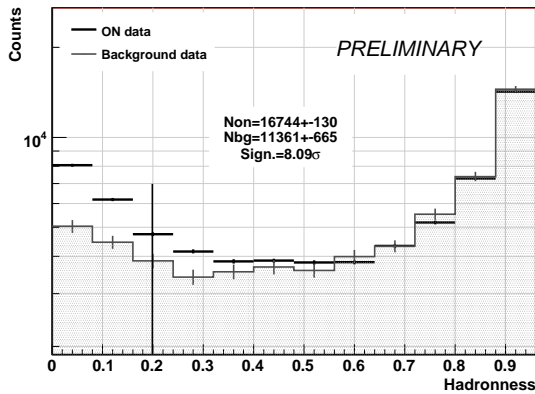


Figure 1: Total *Hadronness* distributions of the ON events (in black) and background simulation (in grey) for energy between 150 GeV and 2 TeV.

Final results and the cosmic electron spectrum will be presented at the conference and included later on in this proceeding.

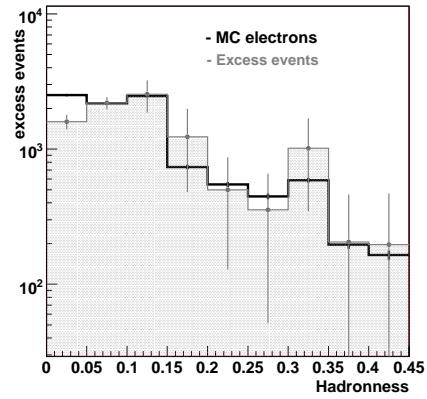


Figure 2: The *Hadronness* distribution of the excess events (in grey) is compared with the one of the MC electrons for energy between 150 GeV and 2 TeV.

## 5 Acknowledgements

We would like to thank the Instituto de Astrofísica de Canarias for the excellent working conditions at the Observatorio del Roque de los Muchachos in La Palma. The support of the German BMBF and MPG, the Italian INFN, the Swiss National Fund SNF, and the Spanish MICINN is gratefully acknowledged. This work was also supported by the Marie Curie program, by the CPAN CSD2007-00042 and MultiDark CSD2009-00064 projects of the Spanish Consolider-Ingenio 2010 programme, by grant DO02-353 of the Bulgarian NSF, by grant 127740 of the Academy of Finland, by the YIP of the Helmholtz Gemeinschaft, by the DFG Cluster of Excellence “Origin and Structure of the Universe”, by the DFG Collaborative Research Centers SFB823/C4 and SFB876/C3, and by the Polish MNiSzW grant 745/N-HESS-MAGIC/2010/0.

## References

- [1] Bi, X.-J., He, X.-G., and Yuan, Q., Phys. Lett., 2009, **B678**: 168-173
- [2] Grasso, D., Profumo, S., Strong, A. W. et al, Astroparticle Physics, 2009, **32**: 140-151
- [3] Blasi, P., Phys. Rev. Lett., 2009, **103**: 051104
- [4] Ackermann, M. et al, Phys. Rev., 2010, **D82**: 092004
- [5] Chang, J. et. al., Nature, 2008, **456**: 362-365
- [6] Aharonian, F. et. al., aap, 2009, **508**: 561-564
- [7] Aharonian, F. et. al., Phys. Rev. Letters, 2008, **101**: 261104-+
- [8] Hillas, M. A., ICRC proceeding, 1985, **3**: 445-448
- [9] Bock, R. K. and Chilingarian, A. and Gaug, M. and Hakl, F. and Hengstebeck, T. and Jiřina, M. and Klaschka, J. and Kotrč, E. and Savický, P. and Towers, S. and Vaiciulis, A. and Wittek, W., Nuclear Instruments and Methods in Physics Research A, 2004, **516**: 511-528

- [10] Maier, G. and Knapp, J., *Astroparticle Physics*, 2007, **28**: 72-81
- [11] Heck, D. et al., *Forschungszentrum Karlsruhe Report*, 1998, **FZKA**: 6019
- [12] CERN-INFN, <http://www.fluka.org>
- [13] Ostapchenko, S. et al., *ICRC proceedings*, 2005, **7**: 135